# The Exploration of the Unknown

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The Square Kilometre Array is conceived as a telescope which will both test fundamental physical laws and transform our current picture of the Universe. However, the scientific challenges outlined in this book are today's problems—will they still be the outstanding problems that will confront astronomers in the period 2020 to 2050 and beyond, when the SKA will be in its most productive years? If history is any example, the excitement of the SKA will not be in the old questions which are answered, but the new questions that will be raised by the new types of observations it will permit. The SKA is a tool for as-yet-unborn users and there is an onus on its designers to allow for the exploration of the unknown. We outline a philosophy for the design and operation of the SKA that can lead the radio astronomers in the 21[st] century to add to the many discoveries of new phenomena made by radio astronomers in the 20[th] century.

## 1. Prologue

> *"Now my own suspicion is that the Universe is not only queerer than we suppose, but queerer than we* CAN *suppose": J. B. S. Haldane*

Most of the phenomena we observe today, using telescopes to observe across the electromagnetic spectrum, were unknown a few decades ago and, to an amazing extent, were discovered by radio astronomers using increasingly powerful instruments and either looking for something else or just following their curiosity. Examples include non-thermal radiation, radio galaxies, quasars, the cosmic microwave background, cosmic evolution, pulsars, gravitational lensing, cosmic masers, molecular clouds, dark matter, and extrasolar planetary systems. These discoveries have changed astronomy in fundamental ways. Some discoveries resulted from increased sensitivity, others from better spatial or temporal resolution, still others by observing in a new wavelength band or even testing misguided theory. Most involved recognizing a new phenomenon and being able to distinguish it from a spurious instrumental response. This scenario is, of course, not restricted to radio astronomy. Perhaps the most spectacular example from astronomy in other wavebands was the discovery of $\gamma$-ray bursts by a military satellite—currently a major field of contemporary astrophysics.

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It is fashionable to imagine that all research follows some classical model of the scientific method—formulation of a model or hypothesis followed by experimental confirmation. Observations not based on testable theoretical predictions are sometimes called "butterfly collecting" or appeals to "serendipity" rather than "real science." Time allocation committees, referees of grant applications, and reviewers of instrument proposals tend to focus on specific questions that will be answered. Yet, astronomy is not an experimental science. We can only observe our Universe and its content with "eyes" as wide-open as possible. We cannot make little changes or experiments to see what happens, except perhaps for some areas of planetary research. So how do we plan for discovery? Despite the apparent capriciousness of our aim, history tells us that a basic requirement is to carry out systematic work with at least an order of magnitude improvement over what has been achieved before in one or more observing capabilities (sensitivity; spatial, temporal, or spectral coverage; spatial, temporal, or spectral resolution). An observing instrument which can offer major advances in several dimensions of parameter space is more likely to make transformational discoveries—history also shows that much greater sensitivity along with flexibility of operation is a wise path to follow. The sensitivity of the Arecibo telescope and the imaging capabilities of array telescopes are excellent paradigms.

Merely providing access to new areas of parameter space with new technology is not a surefire recipe for making ground-breaking discoveries, however. There are other, human, factors to take into account which are just as important for ensuring the SKA's success as a discovery instrument. We are all familiar with what is now the traditional method of using a large common-user telescope involving: i) a proposal to tackle a single small problem; ii) review by time allocation committees; iii) the award of a few hours or maybe even days of observing time; iv) the analysis of the data via a standard suite of software, and, then if all goes well, v) a publication following filtering by a referee. We dub this "the standard model" of observational astronomy and it is perhaps inevitable that the SKA will allocate much of its operations to this analytic mode. When, however, the phenomenon or problem is less well-defined, there can be a rich mix of possible "solutions," only some of which may have been explored by theorists and for which the standard model provides a poor response. It is, therefore, vital to develop a philosophy of design, operations, and data archival which allows individuals, small groups, and larger communities freedom to innovate and encourages users to explore completely new ways of collecting, reducing and analyzing data—in other words *to allow for discovery as well as explanation.*

## 2. The Lessons of Astronomy History

In his 1981 book *Cosmic Discovery* [2] and in subsequent articles, Harwit has addressed the question of what factors lead to new discoveries in astronomy. He argues that a large fraction of the discoveries have been associated with improved coverage of the electromagnetic spectrum or better resolution in the angle, time, or frequency domain. He also notes that astronomical discovery is often closely linked to innovative new technology introduced into the field from outside, often from the military. Consequently, many major new findings have come about more by luck than through careful planning—although what constitutes "luck" is an arguable point that we discuss in §8. Nonetheless theoretical anticipation has usually had little to do with astronomical discovery—what matters most is the implementation of powerful new observing tools.

Will progress at the rate achieved in the second half of the $20^{\text{th}}$ century be likely to continue? Harwit [2] has tackled this seemingly impossible question in two ways. First by estimating the fraction of observational phase space which has presently been explored and then by comparing the number of discoveries attributable to improved instruments with the number independently rediscovered, often by totally unanticipated means, with instruments of quite different kinds. His analysis suggests that we have already seen perhaps 30% to 40% of all the major astrophysical phenomena that will ultimately be revealed by photons, cosmic rays, neutrinos, and captured extraterrestrial material. While one may be sceptical about the quantitative accuracy of this prediction, qualitatively we do not doubt that the Universe still holds plenty of surprises.

In the first half of the $21^{st}$ century, powerful tools in two completely new observational regimes, neutrino and gravitational-wave astronomy, will become available, and it is very likely that they will reveal genuinely new phenomena. Nonetheless, photon astronomy is far from exhausted, and the low energies of radio photons and relative ease with which they are generated and propagate mean that sensitive telescopes in the radio regime will surely contribute their share of new discovery and understanding.

Moreover, radio observations probe a wide range of conditions—from dense gasses to dilute, highly relativistic plasmas—are sensitive to magnetic fields, and yet are not affected by absorption from dust. The fundamental (baryonic) element of the Universe, hydrogen, has a key transition at centimetre wavelengths (the 21-cm hyperfine transition). Radio telescopes routinely make the highest angular resolution observations in astronomy. These capabilities have already been exploited to study some of the most extreme conditions known, e.g., the strong gravitational fields in binary pulsars. It is no surprise that the Key Science Projects currently identified for the SKA exploit all of the above advantages, and we consider it likely that any future discoveries—be they from photons, neutrinos, or gravitational waves—will require radio observations to understand them.

Astronomy at radio wavelengths is marked by a number of differences from shorter wavelength observations, differences that make radio astronomy a powerful technique for observing the sky:

- The sky is mostly empty, which allows unfilled apertures (i.e., interferometers) to operate;
- Long coherent integrations are possible;
- Large numbers of photons are collected so that the signal can be amplified and split without any loss in sensitivity; and
- Diffraction-limited imaging can be obtained via post-processing so that adaptive optics requires no moving parts.

Table 1 lists some of the key discoveries from radio astronomy in the metre and centimetre wavebands and indicates the telescopes and the enabling new parameter space. In addition to adding weight to Harwit's [2] emphasis on the importance of exploiting new technology, several more specific lessons can be learned.

Table 1: Key Discoveries that Illustrate Discovery Space in Radio Astronomy[♯]

| **Discovery** | **Date** | **Enabled By**[♭] | **Telescope** |
|---|---|---|---|
| Cosmic radio emission | 1933 | $\nu$ | Bruce Array (Jansky) |
| Non-thermal cosmic radiation | 1940 | $\nu$ | Reber antennas |
| Solar radio bursts | 1942 | $\nu, \Delta t$ | Radar antennas |
| Extragalactic radio sources | 1949 | $\Delta\theta$ | Australia cliff interferometer |
| 21 cm line of hydrogen | 1951 | theory, $\Delta\nu$ | Harvard horn antenna |
| Mercury & Venus spin rates | 1962,1965 | radar | Arecibo |
| Quasars | 1962 | $\Delta\theta$ | Parkes occultation |
| Cosmic Microwave Background | 1963 | $\Delta S$, calibration, *theory | Bell Labs horn |
| Confirmation of General Relativity (time delay + light bending) | 1964 1970s | theory, radar, $\Delta t$, $\Delta\theta$ | Arecibo, Goldstone, VLA, VLBI |
| Cosmic masers | 1965 | $\Delta\nu$ | UC Berkeley, Haystack |
| Pulsars | 1967 | $\Omega$, $\Delta t$ | Cambridge 1.8 hectare array |
| Superluminal motions in AGN | 1970 | $\Delta\theta$, *theory | Haystack-Goldstone VLBI |
| Interstellar molecules and GMCs | 1970s | theory, $\nu, \Delta\nu$ | NRAO 36-ft |
| Binary neutron stars + gravitational radiation | 1974-present | $\Omega$, $\Delta t$, theory | Arecibo |
| Gravitational lenses | 1979 | $\Delta\theta$, theory | Jodrell Bank interferometer |
| First extrasolar planet system | 1991 | $\Omega$, $\Delta t$ | Arecibo |
| Size of GRB fireball | 1997 | $\lambda\lambda$, $\Delta S$, theory | VLA |

♯ This is a short list covering only metre and centimetre wavelengths.

♭ $\nu \Rightarrow$ spectral coverage; $\Delta\nu \Rightarrow$ spectroscopic resolution; $\Delta S \Rightarrow$ sensitivity; $\Delta t \Rightarrow$ short time resolution.
$\Omega \Rightarrow$ survey with ample sky coverage. $\Delta\theta \Rightarrow$ angular resolution, FoV $\Rightarrow$ field of view, $\lambda\lambda \Rightarrow$ guided by multiwavelength observations;
"theory" $\Rightarrow$ theory played a role in motivating discovery or its search space.
"*theory" $\Rightarrow$ phenomenon was predicted but discovery was independent of the prediction.

- **Discoveries with radio telescopes have set a large part of the current astronomical agenda and radio telescopes are now studying largely what they themselves discovered.**

- **The majority of the discoveries (11/17) were *not* a direct result of theory.** Although there were previous theoretical predictions in two cases, they played no role in the observational discovery.

- **The largest radio telescopes of their day (of a wide range of types) have dominated the discoveries.** This contrasts with Harwit's conclusion that (mainly optical) telescope size was not a major determinant for success. There are several reasons for this difference. Most discrete radio sources are weak, continuum-only emitters. Thus, large radio telescopes, which combine sensitivity and angular resolution, are needed to detect them and to study their characteristics. This contrasts with the situation at optical wavelengths, for which even modest-sized telescopes can observe the myriad of stars with their rich spectroscopic properties. Moreover, the sensitivity of optical telescopes is often limited by photon statistics, so it increases only as the square root of the area of the aperture. For a radio telescope working in the Raleigh-Jeans part of the spectrum, sensitivity scales linearly with aperture. It is notable that there are no filled aperture radio telescopes less than 64 m diameter in great demand at centimetre wavelengths.

- **What a radio telescope was *built for* is almost never what it is *known for*.** Almost invariably, the discoveries in Table 1 were not, often could not have been, in the minds of the designers of those telescopes. For example, Jodrell Bank was built to study meteor trails, Arecibo to study the ionosphere, and the WSRT to do weak source counts. Table 2 shows that the VLA, which is one of the most productive astronomical telescopes of all time, spent only a quarter of its time during its initial decade of operation on the key science drivers listed in the funding proposal.

- **General-purpose telescopes now dominate the discoveries.** While special-purpose instruments dominated discoveries for the first 30 years, the majority of discoveries since then have been made with general purpose telescopes—large filled aperture and arrays of dishes; these are versatile instruments whose performance can be upgraded by new receiving and signal processing capabilities. This trend follows the move to "Big Science" [3] as the cost of facilities with enough sensitivity to continue the exponential growth required for healthy science becomes too expensive for small specialized groups. This lesson encourages us to look for ways in which operational versatility can be built into an inherently common-user instrument like the SKA.

## 3. New Technology for the SKA

It is well established that most scientific advances follow technical innovation. De Solla Price [3] reached this conclusion from his application of quantitative measurement to the progress of science across all disciplines.

Harwit [2] pointed out the most important discoveries in astronomy often result from technical innovation with the discoveries peaking soon after new technology appears, usually within 5 years. However, as a field matures, more general purpose instruments have more impact. During the first 30 years of radio astronomy's brief history, discoveries followed technical innovation, but we are rapidly approaching the limits of obtaining increased capabilities simply by upgrading existing telescopes.

De Solla Price also showed that the normal mode of growth of science is exponential. Historical examples included the rate of discovery of elements and the number of universities founded in Europe. Some more recent examples of exponen-

Table 2
Distribution of VLA Science

| Topic | Observing Time |
|---|---|
| Stars | 16% |
| Galaxies | 14% |
| **Radio Galaxies** | 13% |
| **Quasars** | 9% |
| Star formation | 9% |
| Solar system | 6% |
| AGN | 5% |
| **Cosmology** | 4% |
| Interstellar medium | 4% |
| Supernovae | 4% |
| Galactic Centre | 3% |
| Molecules | 3% |
| VLBI | 3% |
| Pulsars | 2% |
| X-ray, etc. | 1% |
| Astrometry | 1% |

Items in bold were key scientific drivers in 1967 funding proposals. List compiled by RDE during his tenure at the VLA.

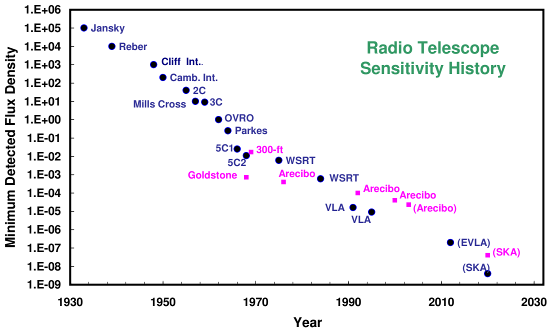


Figure 1. The history of the weakest radio objects detected using radio telescopes in both imaging (radio sources, circles) and time-domain (pulsars, squares) experiments. The points for the EVLA and the SKA are projected based on the expected thermal noise limit for a 100-hour integration.

tial growth are particle accelerator beam energy, Internet hosts, and the now famous "Moore's Law" for computing devices. Such exponential growth cannot continue indefinitely without a re-organisation or change in technology.

Figure 1 shows the weakest detected radio sources as a function of time from the start of radio astronomy before World War II to the present day and beyond. Exponentially increasing sensitivities have produced an improvement in the weakest sources detected, for both continuum sources and pulsars, by a factor of roughly $10^6$ over the past 60 years. These exponential improvements in the detection of weak sources have come from a series of technical innovations involving a combination of increasing collecting area, decreasing noise temperature, increasing bandwidth, and development of algorithms for wide-field, high-dynamic range imaging.

Can we continue the exponential growth? Do we have a new technology? Receivers are operating near their fundamental sensitivity limits, bandwidths comparable to the observing frequency will soon be achieved, time-bandwidth products of order unity are being analyzed, and interferometry has exploited antenna separations comparable to the Earth's diameter. For radio astronomy, the technical challenge is now to obtain increased collecting area at a reasonable cost, particularly so for observations of spectral line emission where an increased bandwidth has no effect on the sensitivity. Thus, the original arguments for a square kilometre aperture SKA based on detecting and imaging the 21-cm line radiation of atomic hydrogen at high redshift remains.

The combination of transistor amplifiers and their large scale integration into complex systems which can be duplicated inexpensively provides one of the keys for change. The other key technology is the computing capacity to apply digital processing at high bandwidth, thereby realizing processes such as multiple adaptive beam formation and active interference rejection in ways not previously conceivable. The SKA is very demanding of new technology but the SKA vision has unleashed an unprecedented burst of creativity in the radio astronomy and engineering communities.

Technologies that must be developed to realize the SKA include i) wider bandwidth radio frequency receivers and their direct integration into antennas; ii) low-cost manufacturing of large collector systems; iii) low-cost signal digitization and conditioning; iv) radio frequency interference (RFI) mitigation techniques; v) low-cost computing for real-time beam-forming and data processing; vi) wide-band data transmission via optical fibres; vii) high speed real-time imaging of massive data sets, and viii) applying emerging computational tools for data mining to the huge data sets that will be generated. In contrast to the previous and current generation of receiver and signal processing equipment, which uses special-purpose hardware, the next generation will inevitably exploit the convergence of radio and digital computing technologies—replacing hardware with firmware or software and allowing unprecedented versatility via the use of programmable processing engines.

More than simply a quantitative increase in collecting area, though, the application of these new technologies promises qualitatively new avenues for observing the sky. The SKA has the potential to be a highly multiplexed instrument, involving multi-beaming on a much larger scale than can be envisioned with current telescopes, so that multiple users can have access to the telescope simultaneously for independent projects. As computing power and memory become cheaper, the SKA could be "digitally upgraded," evolving into a more powerful telescope with time. Exploiting the possibility of rapid and large disk and memory (RAM) storage of digital data, the SKA has the potential to provide a "look-back" buffer, allowing the telescope to be pointed in an *ex post facto* manner based on a trigger, say, from another telescope at another wavelength. Finally, its digital underpinnings naturally lead to rich contributions to Virtual Observatories of the future.

## 4. SKA Parameter Space

The SKA cannot explore much new parameter space because angular resolution, spectral-temporal resolution, and polarization parameters have all been probed at most wavelengths with the current generation of radio telescopes. However, the SKA will greatly enlarge known parameter space by i) a much greater sensitivity, over a very wide range of angular resolutions; ii) a much larger instantaneous field-of-view at least at some wavelengths; iii) the potential for multiple, independently steerable, fields-of-view within which are independently steerable beams. The sensitivity and sky coverage advances combine to provide two major steps forward compared with current instruments i) the volume of space accessible to the SKA will be enormously increased, and hence the chances of finding intrinsically rare objects in large scale surveys will be much enhanced; ii) the potential for "all-sky" coverage allows for statistical analysis of the surveys not biased by the small area of sky being studied. A major advantage of multiple independent beams is a multiplex one: several groups pursuing different goals can operate on the SKA simultaneously and "light" will thereby stream in through more windows on the Universe.

One intriguing expansion into new parameter space is negative time! A time buffer recording raw data (undetected voltages) would allow an astronomer (or more likely a computer) to interrogate the raw SKA data and form a full sensitivity beam anywhere in the field of view (FoV) for periods of ten seconds to minutes earlier in time, hence the SKA can have started monitoring *before* receipt of a trigger. This trigger could be observations by the SKA itself or from some other instrument, such as a satellite monitoring $\gamma$-ray bursts.

The second fundamental design driver of the SKA is the instantaneous FoV. Because radio wavelengths are long, a large FoV is natural but changing technology may enable even greater increases. The SKA design goal has specified an FoV of at least 1 deg$^2$ at a wavelength of 21 cm. In some concepts being pursued, this may rise to as much as 100 deg$^2$ or more.

While the instantaneous FoV that might be achieved for the SKA is impressive ($> 10$ deg$^2$) when compared to modern radio telescopes, it is smaller (in some cases much smaller) than some historical radio telescopes. The SKA's potential to make a qualitative improvement over previous and existing telescopes is its simultaneous combination of large instantaneous FoV with high sensitivity and angular resolution. Historically, radio telescopes have been capable of obtaining either large solid angle coverage (e.g., the STARE survey at 610 MHz by Katz et al. [4] with a FWHM beam of $4000'$) or high sensitivity (e.g., the Arecibo telescope with a gain of 11 K $\mathrm{Jy}^{-1}$ at 430 MHz) but not the two simultaneously. If the SKA satisfies only its design requirements, it will produce at least an order of magnitude sensitivity improvement, over a large frequency range, with a solid angle coverage that is comparable to or exceeds that of all but a small number of low-sensitivity telescopes.

The SKA's potential is even more striking given that all of the Key Science Projects identified for the SKA project involve surveys of one sort or another. The surveying speed of a telescope system to reach a given flux density limit is proportional to the product of its instantaneous FoV and the square of its sensitivity. In Table 3 we compare the relative continuum surveying speeds of current instruments with that of the SKA. It is clear that the SKA will have a surveying speed at least 10,000 times greater than is possible with the current instruments. If the advantages of the wide FoV concepts being investigated can be realised, then a factor larger than $10^5$ can be achieved. New discoveries often arise from large-scale surveys during which the rare discrete objects of a new type or new large-scale emergent properties of the ensemble of objects stand out. The advance offered by the SKA promises a complete revolution in our ability to survey the radio sky—so far in advance of anything which has been done before that it truly can claim to take the SKA into an arena of "discovery science."

## 5. New Opportunities for the SKA

Although the SKA will not investigate new dimensions in parameter space per se, its capabilities will allow new opportunities for investigating the radio sky. Here we summarize two of the avenues that seem most promising today, recognizing that other possibilities may arise in the next 20–50 years.

Table 3
Relative Surveying Speed for Continuum Sources

| Telescope | Relative Sensitivity (per beam) | Field of View ($\mathrm{deg}^2$) | Survey Speed |
|---|---|---|---|
| SKA specification ($A_{\mathrm{eff}}/T_{\mathrm{sys}} = 20,000$) | 1 | 1 | 1 |
| SKA potential | 1 | 100 | 100 |
| Arecibo/ALFA ($A_{\mathrm{eff}}/T_{\mathrm{sys}} = 760$) | 0.038 | 0.03 | $10^{-4.4}$ |
| VLA ($A_{\mathrm{eff}}/T_{\mathrm{sys}} = 220$) | 0.011 | 0.26 | $10^{-4.5}$ |
| Parkes ($A_{\mathrm{eff}}/T_{\mathrm{sys}} = 80$) | 0.004 | 0.52 | $10^{-5.1}$ |
| GBT/Effelsberg ($A_{\mathrm{eff}}/T_{\mathrm{sys}} = 240$) | 0.012 | 0.02 | $10^{-5.7}$ |

The tabulated surveying speeds are applicable to observations at a wavelength of 21 cm. The surveying speed is calculated as the square of the sensitivity multiplied by the instantaneous FoV. In calculating the relative sensitivity per beam, all continuum bandwidths have been assumed to be equal.

### 5.1. The Dynamic Radio Sky

A combination of high time resolution, high sensitivity, and large instantaneous FoV will enable the SKA to search for and study transient sources. Although various classes of radio transients are known today—such as giant pulses from radio pulsars, radio flares from micro-quasars and brown dwarfs, maser flares, supernovae events, and radio afterglows from $\gamma$-ray bursts—except for pulsars, there have been few surveys of a large area of sky at radio wavelengths for transient sources.

The success of X- and $\gamma$-ray telescopes in carrying out such wide-field surveys suggests that wide-field radio surveys would be equally fruitful. Figure 2 illustrates the potential range of radio transient phenomena is large. In the Rayleigh-Jeans approximation, a source with brightness temperature $T$ varies (intrinsically) on a time scale or pulse width $W$,

$$W^2 = \frac{1}{2\pi k}\frac{SD^2}{T}\frac{1}{\nu^2}, \qquad (1)$$

where the observed flux density is $S$, the source's distance is $D$, the emission frequency is $\nu$, and $k$ is Boltzmann's constant. As Figure 2 shows, the range of $\nu W$ covers at least 13 orders of magnitude while the range of $SD^2$ covers at least 20 orders of magnitude. Moreover, in addition to the known classes of radio transients, a variety of other classes involving either coherent or incoherent processes have been suggested such as

- Radio flares from extra-solar giant planets, akin to Jovian decametric radiation;
- Prompt emission from supernovae, as in SN 1987A;
- Radio-loud, $\gamma$-ray quiet "$\gamma$-ray bursts";
- Prompt emission from $\gamma$-ray bursts;
- Radio hypernovae associated with Population III stars;
- Galactic flares, outbursts due to the tidal disruption of single stars falling into the supermassive black holes at the centre of apparently normal galaxies;
- Evaporating black holes; and
- Extraterrestrial transmitters.

### 5.2. Synergy with Other Telescopes

The sensitivity of the SKA means that for the first time many objects which dominate the science done at other wavebands (stars, galaxies, star forming regions, proto-galaxies, galaxy clusters) become readily observed at radio wavelengths. The SKA will be a leading member of a complementary group of next-generation photon collectors including: ground-based optical telescopes in the 30–50 m class; the JWST (near- and mid-IR); ALMA (mm- and sub-mm wave); next-generation X-ray and $\gamma$-ray observatories (XEUS and Constellation-X). All of these will provide imaging information on the scale of roughly 0.1 arcsecond or better and all will provide unique views of the Universe. Completely new information will then flow from a comparison of the enormous data sets which will be produced from synoptic surveys in many wavebands. From a combination of these surveys, statistically significant patterns and subtle correlations between parameters will become apparent, pointing to new phenomena and rare or previously unknown objects will stand out. Meanwhile the first generations of neutrino detectors, the ground-based gravitational wave observatories (LIGO; GEO-600; VIRGO, and others) and the space-based follow-ons (LISA) are expected to lead to the detection of gravitational waves from a variety of sources. All of these ventures and their anticipated outcomes in terms of discovery and understanding will benefit from observations in the radio band. The SKA will provide the sensitivity along with a rich set of observational modes to enable exploration of the radio sky in all of its anticipated complexity.

## 6. Design Philosophies for the SKA

In designing the SKA, we need to look beyond an instrument which will merely satisfy the requirements of the current Key Science Projects—challenging though these are—and envisage an instrument which also allows new observing paradigms. Some guiding principles are

- Seek the maximum flexibility in the design, to enable the combination of high sensitivity simultaneously with high resolution in time, frequency and/or angle. Where multiple requirements cannot be met simultaneously, allow the ability to make trade off between different aspects, e.g., collecting area against solid-angle coverage, or signal processing bandwidth at high frequencies against FoV at low frequencies.
- Seek ways for the multiple reuse of expensive components, for example via simultaneous multiple beams or spectral bands.
- Exploit the areas of fastest changes in technology by designing the signal processing hardware in such a way that the major purchases can be made late in the project.
- Plan the architecture for upgrades right from the start by adopting an open modular

approach to the various systems and subsystems. In particular digital signal processing and computing power will increase by many orders of magnitude while decreasing in cost over the course of the SKA's lifetime. This will allow us to generate more beams and hence more channels through which to examine the sky. The SKA's architecture should recognize this and provide clear interfaces, down-stream of which it is relatively easy to upgrade the electronics in a step-wise fashion.

- Begin the software development early; the SKA will inevitably be the largest software project in astronomy, and there is no Moore's Law escalator to help.

- Facilitate time-buffering and archiving as much of the raw data as possible to enable the pre-trigger beam-forming mode discussed in §3

Ideally the SKA will be designed so as to provide independently-steerable fields-of-view and independently-steerable beams within each FoV, over at least the lower portion of its operational frequency range. Such architecture would provide users with a highly flexible and responsive instrument and would inevitably lead to changes in observing style. It remains to be seen how feasible it is to achieve all this flexibility from the start. However, we note that such an approach generically incorporates

- A science survey advantage, which is required for a range of science programmes requiring huge amounts of telescope time and which would be impossible with conventional systems;

- Long term monitoring opportunities with dedicated beams;

- A "community" advantage because many groups could access the whole aperture simultaneously, allowing the operation of the SKA to resemble that of particle accelerators or synchrotron light sources;

- A multiplex advantage, simply by increasing the range of different data sets which can be collected; and

- An adaptive beam-forming advantage, as "reception nulls" could be steered to cancel out sources of RFI.

In addition to the Key Science Projects, the different FoVs could, for example, be used for

- Monitoring variable sources such as looking for pulsar glitches and timing changes due to gravitational radiation or observing effects of interstellar scintillation;

- Integrating for long periods of time to obtain the ultimate in sensitivity;

- Studies of time variable phenomena on time scales from nanoseconds and longer and seeking transient radio sources and responding instantly to transients discovered in other wavebands; or

- Experimentation or "high risk" observations which would not be scheduled by the standard peer-review process.

## 7. Exploration-Driven Observing

The SKA will be sufficiently different from current instruments that the "standard model," discussed in §1, via which the user interacts with the instrument and its data, can be re-evaluated. It is vital that the flexible modes of use of the SKA not become too formalised. Many astronomers are conscious that the standard model is not the way to make breakthroughs and other responses are already being tried. The NVSS and FIRST surveys and the Hubble Deep Field are examples of one different approach where some of the time on a telescope is devoted to large programs whose data products are then made available publicly. The success of these programmes encourages us to look for further ways to break the constraints of the traditional common-user paradigm and an excess of egalitarianism.

Even the existing science goals for the SKA demand a great deal of flexibility in the way the total collecting area, frequency coverage, and angular resolution can be exploited. For example, some observations require high resolution in time, frequency, and/or angle along with the ability to trade off collecting area against solid-angle coverage. For the SKA we wish to use the sensitivity for both imaging and non-imaging science in ways that adopt the best features of existing arrays and filled aperture antennas.

Modes of operation carry with them implications for data processing and analysis. Data rates are unprecedented and will require large increases in computing capacity and new algorithms over the next 10–15 years and beyond. Brute force solutions (all of everything) will not be practical so trade-offs will have to be made. Flexible implementations of the various modes will allow these trade-offs between resources while maximizing the scientific value of the observations, but when searching for the unknown we must avoid making the observational filter so narrow that we see only what we already know to exist.

Particular modes of observation that need to be enabled in order to address the Key Science Projects and especially for the exploration of as-yet unknown phenomena include

**Full Field-of-view Imaging** This mode will be the one used by most of the surveys in the Key Science Projects. Enabling of this mode requires channel widths and correlator dump times that scale as the reciprocal of the maximum baseline. In order to avoid excessive signal transport and processing loads, this mode would normally be limited to the central part of the array.

**High-resolution Imaging** With the SKA, real time correlation for high-resolution imaging will be possible with data transport by fibre. However, data rates will not allow full FoV imaging using the longest baselines in the envisioned array. Smaller regions can be processed at the full resolution, and it should be also possible to image multiple smaller patches, e.g., around each radio source in the FoV.

**Planetary Radar/Time-Gated Imaging** High resolution imaging of radar return signals can remove the ambiguities usually associated with delay-Doppler radio imaging can be removed by forming delay-Doppler images within each resolution element of a synthesized image. Similar gating can be used to analyse other periodic signals, e.g., to "gate out" pulsars in a search for un-pulsed emission.

**Slow Transient Searches** Slow transients are defined as those with durations long enough that they can be found or studied by moving the FoV (e.g., a raster scan) over the region being studied. For GRB afterglows, for example, the entire available sky could be sampled in 1 day using a mosaic imaging mode that provides roughly 2 s of integration per direction even for a 1 deg$^2$ FoV.

**Fast Transient Searches** Fast transients require a long-dwell staring mode. For wide-field blind surveys, the instantaneous solid angle needs to be maximized. This places a limit on maximum baseline that can be used in order to keep data processing rates manageable. The fastest transients, e.g., $\Delta t \leq 1$ s, are subject to distortion by dispersive propagation like that seen from pulsars and therefore require adequate frequency resolution for "dedispersion" methods to be applied.

**Blind Pulsar Surveys** For pulsars, which are essentially a special case of fast transients, it will be possible to analyse only the inner core array of size $b_{\max} \sim 1$ to 5 km. Even so, this already implies about $10^4$–$10^5$ pixels and the analysis for each pixel includes summing over frequency using roughly $10^3$ trial dispersion measure values, followed by Fourier analysis of each resulting time series and statistical threshold tests. In order to find pulsars in compact binaries with orbital periods less than a few hours, a search over an acceleration parameter must also be done for a few hundred values.

**Blind SETI Surveys** SETI usually involves searches for possibly modulated, very narrow-band carrier signals. Interstellar propagation limits the minimum signal bandwidth to about 0.05 Hz for distances larger than about 300 pc at 1 GHz. Spectra must therefore be computed with billions of channels and concordant time resolution for each pixel. As with pulsars, blind surveys will use the inner core array.

**Targeted Observations** When a source's location is known or constrained from other astronomical observations, processing requirements are lessened significantly and can exploit real-time beamforming capabilities. With multiple beams within the primary FoV, searches can be made on multiple targets. Analyses of signals from each beam may require special-purpose hardware (SETI spectrometers, pulsar-timing machines, etc.) to handle the data rates.

**Incoherent Summing Modes** When the full gain of the SKA is not needed, but access to the entire primary FoV is desired, signals from antennas can be summed incoherently. For $N$ antennas divided into $N_{\rm sa}$ subarrays, the sensitivity per analyzed subarray signal is $(G/T)_{\rm sa} = (G/T)_{\rm SKA}/(N \times N_{\rm sa})^{1/2}$. With this mode, for example, 100 subarrays could yield instantaneous coverage of 100 deg$^2$ and the sensitivity equivalent of a 30-m antenna. This mode can provide very high time resolution without requiring a very high data processing rate and may find use with the SKA early on in its deployment.

**RFI Mitigation Modes** Data rates are also entwined with issues pertaining to shared use of the radio spectrum, which is likely to grow worse with time. Specialized data acquisition modes for pulsars and transients, with high resolution in time and frequency, may also be necessary to ensure that RFI mitigation algorithms, which might otherwise assume that interfering signals fill only fractions of the overall time-frequency plane, do not excise the very signal being sought. Clearly sites with low levels of terrestrial interference are preferred in order to maximise the probabilities of recognising unexpected types of signals.

## 8. The Human Factor

*"In the field of observation, chance favours the prepared mind": L. Pasteur*

*"The harder I practice, the luckier I get": Sports adage*

So far we have addressed essentially technological solutions to the question, How do we make the SKA perform as a discovery instrument? These can all be gathered together within the generic answer, Provide access to new regions of a multiparameter space via the application of innovative radio technology. There is more to discovery than can be plotted as parameter space on a graph, though—there is an abstract or innovation aspect to discovery as well, which is just as important but is often overlooked. We should be inspired by another lesson of history.

*Astronomical discoveries are usually made by people who are "curious" and take the time to understand their instrument, with less emphasis on a quick publication.*

The lesson is to allow people enough "room" to make discoveries. Discoveries invariably result from an individual becoming so familiar with the data, *and hence the possible sources of error in them*, that he/she can recognize and unexpected clue for what it is worth. The discovery of pulsars is a perfect example. It was only the apparently tedious day-to-day routine of checking miles of pen-recorder charts that enabled Jocelyn Bell to spot her famous "bits of scruff" and to distinguish them from the usual types of interference. She paid her dues and for the right mind familiarity does not breed contempt for the data. The discovery of Jupiter bursts, interplanetary scintillations, the CMB, and even the original discovery of Jansky himself, are other examples. While we cannot control the subtleties of individual human curiosity, we can provide the circumstances in which it can flourish.

The continued success stories of pulsar groups around the world, stretching over nearly four decades in time, provides further lessons for SKA planners. In addition to choosing a field which is phenomenologically rich and for which the inverse problem, connecting data to fundamental physics, is tractable, what else have they done right?

- Designed their own "back-end" equipment and software and constantly updated it.
- Archived the raw data and re-analysed them as computing power increased.

Some possible responses to these lessons for SKA planners are

- Award some time to successful groups or collaborations on the basis of their past record—a "rolling time allocation grant" which can be sustained or closed down on the basis of performance integrated over several years.
- Allow high-risk or unproven new-style observations. The availability of independent beams will help to make this feasible without compromising conventional observing programs.
- Build an open system that allows user-produced hardware and software to be employed on the telescope. Clean interfaces will need to be defined and maintained.
- Maintain the technical expertise in community. The experience in the USA over the past two decades during which technology development in many of the traditional university departments systematically fell away as resources were directed to the national facilities, should act as a warning. In the lead up to the SKA innovative R&D is re-invigorating the world-wide community and involving a new generation of engineers and students. When the SKA is completed, it is vital to allow a cadre of technical people world-wide to gain continuous access to parts of the system for continuous experimentation.

## 9. Summary

The impact of the SKA will depend not so much on the cleverness of the astronomers in defining the important astrophysical problems of today, but in the cleverness and ambition of its designers to obtain better sensitivity, higher dynamic range, larger field-of-view, and perhaps some other parameters not yet contemplated.

The SKA planners and designers must therefore have a vision for a system that is not only much more powerful than previous radio telescopes but at the same time is highly flexible, easy-to-use, and has an operating philosophy which positively encourages and allows the astronomers of tomorrow to look at the sky and to examine their data in new and creative ways. It must also allow the radio engineers of tomorrow the space to conceive and design the innovative upgrades which increased signal processing and computing power will allow.

Just as important will be to attract and nourish talented and careful observers who can become so familiar with particular types of data as to be able to spot the important new clues which will lead to fresh discoveries. The SKA will need to solve their problems, not our problems, which they will have already solved or shown to be naive or irrelevant. Including "Exploration of the Unknown" as one of the prime goals of a highly-flexible SKA is founded firmly on the profound contributions which radio astronomy has already made to our knowledge of the Universe.

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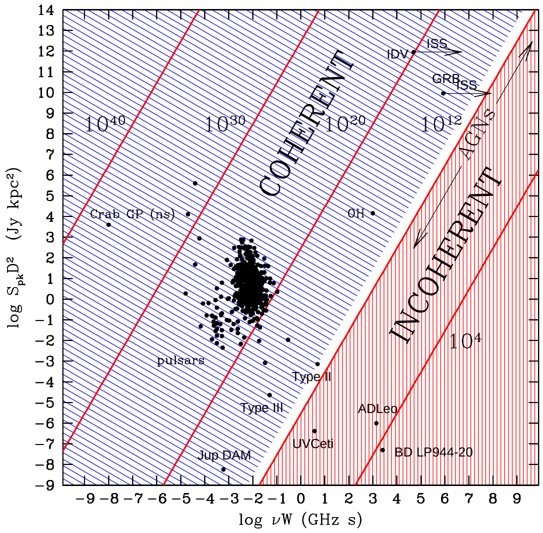


Figure 2. Phase space for known and anticipated transient signals. The horizontal axis is the product of transient duration $W$ and radio frequency $\nu$ while the vertical axis is the product of the peak flux density $S_{\rm pk}$ and the square of the distance $D$ and is proportional to luminosity. Lines of constant brightness temperature are shown, and the uncertainty principle requires that signals be to the right of $\nu W = 10^{-9}$. Also shown are representative examples of active galactic nuclei (AGNs); $\gamma$-ray burst afterglows (GRBs); pulsars, including the 2 ns giant pulses from the Crab pulsar (Crab GP [ns], [1]); flare stars and brown dwarfs; masers; planets; and the Sun. In addition, propagation effects produce intraday variability (IDV) and interstellar scintillation (ISS), which modulate compact radio sources on a variety of time scales. These effects make sources appear to be more variable than they are intrinsically but are rich in information about the intervening media. The known objects span nanosecond to year durations. The dynamic radio sky represents an opportunity for the SKA to expand our knowledge of the radio sky in the same way that high-energy transient studies, conducted with wide-field X- and $\gamma$-ray instruments, have discovered $\gamma$-ray bursts, accreting black holes and other compact objects, and flares from planets, stars and AGNs.